Neutrinos, where BSM physics begins (III)

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Neutrino Frontiers, GGI
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HIDDeN
Hunting Invisibles: Dark sectors, Dark matter and Neutrinos

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Non standard neutrino interactions

They can be described by effective four-fermion operators of the form

$$2\sqrt{2}G_F\varepsilon_{\alpha\beta}\left(\bar{\nu}_\beta\gamma^\mu P_L l_\alpha\right)\left(\bar{f}\gamma_\mu P_{L,R} f'\right)$$

normalizing the operator with the Fermi constant

$$\varepsilon_{\alpha\beta} = \frac{M_W^2}{M_{NSNI}^2}$$

NSNI can appear at every step. It is therefore necessary to break down the analysis in three stages

- the production process
- the time evolution
- the detection process

We are left "only" with neutral current NSNI

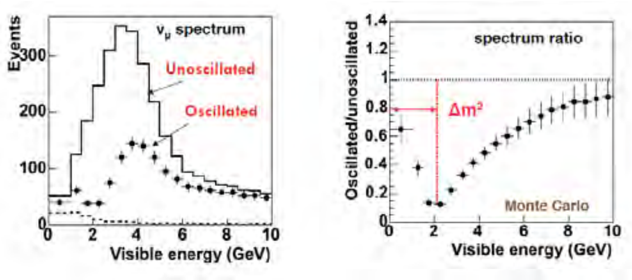
$$2\sqrt{2}G_F\varepsilon_{\alpha\beta}\left(\overline{\nu}_\beta\gamma^\mu P_L\nu_\alpha\right)\left(\bar{f}\gamma_\mu P_{L,R}f\right)$$

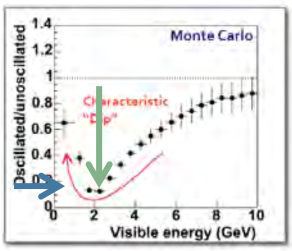
$$i\frac{d}{dt}\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix}=\frac{1}{2E}\left[U\begin{pmatrix}0&0&0\\0&\Delta m_{21}^2&0\\0&0&\Delta m_{31}^2\end{pmatrix}U^\dagger+a\begin{pmatrix}1+\varepsilon_{ee}&\varepsilon_{e\mu}&\varepsilon_{e\tau}\\ \varepsilon_{e\mu}^*&\varepsilon_{\mu\mu}&\varepsilon_{\mu\tau}\\ \varepsilon_{e\tau}^*&\varepsilon_{\mu\tau}^*&\varepsilon_{\tau\tau}\end{pmatrix}\right]\begin{pmatrix}\nu_e\\ \nu_\mu\\ \nu_\tau\end{pmatrix}$$

$$a = 2\sqrt{2}G_F n_e E$$

$$H=\frac{1}{2E}\left[U\begin{pmatrix}0&\\&\Delta m_{32}^2\end{pmatrix}U^\dagger+a\begin{pmatrix}\epsilon_{\mu\mu}&\epsilon_{\mu\tau}\\ \epsilon_{\mu\tau}^*&\epsilon_{\tau\tau}\end{pmatrix}\right]$$

$$P(\nu_\mu \to \nu_\mu) = 1 - \sin^2(2\theta)\sin^2(1.27\Delta m^2 L/E)$$





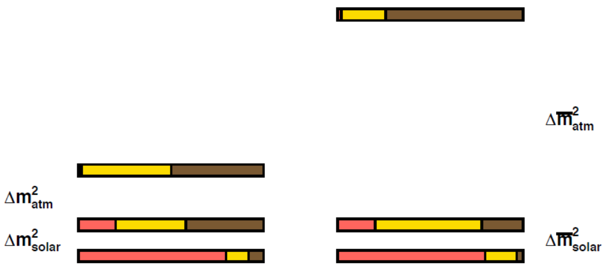
$\varepsilon_{\mu\tau}$ changes the disspearence probability at large energies shifts the position of the minimum in energy

Δm^2

$\varepsilon_{\tau\tau}$ modifies the dissapearence probability near the first oscillation minimum, especially the depth of the minimum

$sin^2(2\ \theta_{23})$

CPT violation



$$\frac{|\ m(K_0)\ - m(\overline{K_0}\,)\ |}{m_{K-av}} < 10^{-18}$$

$$m_{K-av} \approx \frac{1}{2}\ 10^9\ \text{eV}$$

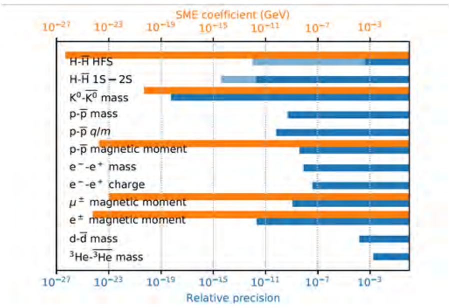
$$(m(K_0) - m(\overline{K_0}\,)\ )(m(K_0) + m(\overline{K_0}\,)) < 2\ \ 10^{-18}\, {m_{K-av}}^2$$

$$|m^2(K_0)\ - m^2(\overline{K_0}\,)| \approx \frac{1}{2}\ \text{eV}^{\,2}$$

CPT tests

CPT invariance tested in several matter-antimatter systems:

- neutral kaons
- electron/positron
- proton/antiproton
- H/anti-H



E. Widmann, arXiv:2111.04056 [hep-ex]

Several experiments at the Antiproton Decelerator and ELENA(Extra Low Energy Antiproton) @CERN

Current bounds

- We can use data of various experiments to calculate the neutrino and antineutrino oscillation parameters:
  - Solar neutrino data: $\theta_{12}, \Delta m^2_{21}, \theta_{13}$
  - Neutrino mode in LBL: $\theta_{23}, \Delta m^2_{31}, \theta_{13}$
  - KamLAND data: $\overline{\theta}_{12}, \Delta\overline{m}^2_{21}, \overline{\theta}_{13}$
  - SBL reactors: $\overline{\theta}_{13}, \Delta\overline{m}^2_{31}$
  - Antineutrino mode in LBL: $\overline{\theta}_{23}, \Delta\overline{m}^2_{31}, \overline{\theta}_{13}$
- No bounds on CP-phases since all values are allowed

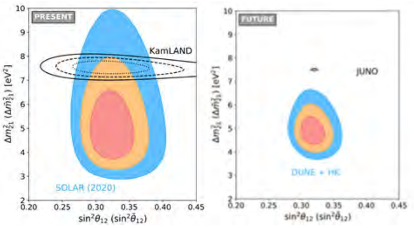
Parameter	Main contribution	Other contributions
θ_{12}	SOL	KamLAND
θ_{13}	REAC	ATM+LBL and SOL+KamLAND
θ_{23}	ATM+LBL	-
δ_{CP}	LBL	ATM
Δm^2_{21}	KamLAND	SOL
$\lvert\Delta m^2_{31}\rvert$	LBL+ATM+REAC	-
MO	LBL+REAC and ATM	-

SOL: Solar
ATM: Amtopsheric neutrinos
LBL: Long baseline accelerator experiments
REAC: Short-baseline reactor experiments

Current bounds

- We use the same data (except atmospheric neutrinos) as for the global fit to obtain

$$
\begin{aligned}
|\Delta m_{21}^2 - \Delta \overline{m}_{21}^2| &< 4.7 \times 10^{-5}\ \mathrm{eV}^2, \\
|\Delta m_{31}^2 - \Delta \overline{m}_{31}^2| &< 2.5 \times 10^{-4}\ \mathrm{eV}^2, \\
|\sin^2\theta_{12} - \sin^2\overline{\theta}_{12}| &< 0.14, \\
|\sin^2\theta_{13} - \sin^2\overline{\theta}_{13}| &< 0.029, \\
|\sin^2\theta_{23} - \sin^2\overline{\theta}_{23}| &< 0.19.
\end{aligned}
$$

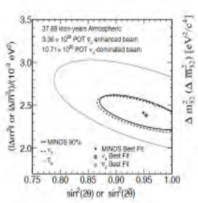


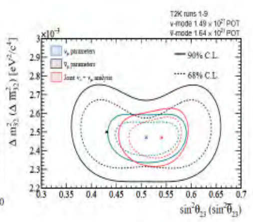
T2K results, a hint ?

- T2K studied neutrino and anti-neutrino oscillations separated

$$\sin^2\theta_{23} = 0.51, \quad \Delta m_{32}^2 = 2.53 \times 10^{-3}\mathrm{eV}^2$$
$$\sin^2\overline{\theta}_{23} = 0.42, \quad \Delta \overline{m}_{32}^2 = 2.55 \times 10^{-3}\mathrm{eV}^2$$

- Results are consistent with
- CPT-conservation





- In experiments and in fits normally you assume CPT-conservation
- If CPT is not conserved this leads to impostor (fake) solutions in the fits
- To perform the standard fit you would calculate

$$\chi^2_{\mathrm{total}} = \chi^2(\nu) + \chi^2(\overline{\nu})$$

and then minimize this function

$$h(x,y) = f(x) + g(y)$$

$$\partial_x f(x) = 0 \qquad \partial_y g(y) = 0$$

$$x = y$$

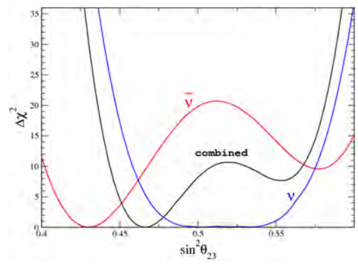
$$h(x) = f(x) + g(x)$$

$$\partial_x f(x) = \partial_x g(x) = 0$$

$$\partial_x f(x) = -\partial_x g(x)$$

Obtaining impostor solutions

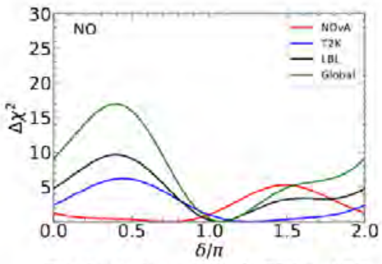
- This was done for $\sin^2(\theta_{23}) = 0.5,\ \sin^2(\overline{\theta}_{23}) = 0.43$



Combined best fit value is now $\sin^2(\theta_{23}^{\mathrm{comb}}) = 0.467$

Real true values are disfavored at close to 3σ and more 5σ confidence levels

$\theta_{13} \neq \theta_{13}$ can account for different behavior in neutrino and antineutrino channels



Tension between NOvA, T2K and SK atm. and $\delta_{bf} = 1.08\pi$

- Disfavours:
  - $\delta = \pi/2$ at 4.0σ
  - $\delta = 0$ at 3.0σ
  - $\delta = 3\pi/2$ with $\Delta\chi2 = 4.9$

all values of δ and $\overline{\delta}$ remain allowed at ~ 1σ

The increasing precision in neutrino oscillation measurements requires a thorough analysis of the assumptions considered.

Distinguishing CPT violation from NSNI

The muon neutrino survival probability in matter can be written as

$$P(\nu_\mu \to \nu_\mu) = 1 - \sin^2 2\theta_\nu \sin^2\left(\frac{\Delta m_\nu^2 L}{4E}\right).$$

in matter

$$\begin{aligned} \Delta m_\nu^2 \cos 2\theta_\nu &= \Delta m^2 \cos 2\theta + \epsilon_{\tau\tau} A\,, & \Delta m_{\bar\nu}^2 \cos 2\theta_{\bar\nu} &= \Delta m^2 \cos 2\theta - \epsilon_{\tau\tau} A\,, \\ \Delta m_\nu^2 \sin 2\theta_\nu &= \Delta m^2 \sin 2\theta + 2\epsilon_{\mu\tau} A\,. & \Delta m_{\bar\nu}^2 \sin 2\theta_\nu &= \Delta m^2 \sin 2\theta - 2\epsilon_{\mu\tau} A\,. \end{aligned}$$

$$4\Delta m^4 = \Delta m_\nu^4 + \Delta m_{\bar\nu}^4 + 2\Delta m_\nu^2 \Delta m_{\bar\nu}^2 \cos(2\theta_\nu - 2\theta_{\bar\nu})$$

$$\sin^2(2\theta) = \frac{(\Delta m_\nu^2 \sin(2\theta_\nu) + \Delta m_{\bar\nu}^2 \sin(2\theta_{\bar\nu}))^2}{\Delta m_\nu^4 + \Delta m_{\bar\nu}^4 + 2\Delta m_\nu^2 \Delta m_{\bar\nu}^2 \cos(2\theta_\nu - 2\theta_{\bar\nu})}$$

$$\begin{aligned} 2\epsilon_{\tau\tau}^m A &= \Delta m_{\bar\nu}^2 \cos(2\theta_\nu) - \Delta m_{\bar\nu}^2 \cos(2\theta_{\bar\nu}) \\ 4\epsilon_{\mu\tau}^m A &= \Delta m_\nu^2 \sin(2\theta_\nu) - \Delta m_{\bar\nu}^2 \sin(2\theta_{\bar\nu}) \end{aligned}$$

G.B., C. Ternes and M. Tortola, Eur.Phys.J.C 79 (2019) 5, 390

Violations of Lorentz invariance

Lorentz violation

$$(h_{\text{eff}})_{ab} = \frac{m^2_{ab}}{2E} + \frac{1}{E}[(a_L)^\alpha p_\alpha - (c_L)^{\alpha\beta} p_\alpha p_\beta]_{ab}$$

standard Lorentz covariant term

violates both CPT and Lorentz invariance

Violations of Lorentz invariance

Lorentz violation

$$(h_{\text{eff}})_{ab} = \frac{m^2_{ab}}{2E} + \frac{1}{E}[(a_L)^\alpha p_\alpha - (c_L)^{\alpha\beta} p_\alpha p_\beta]_{ab}$$

standard Lorentz covariant term

violates both CPT and Lorentz invariance

As usual, the oscillation probability is governed by the difference of the eigenvalues of the effective hamiltonian.

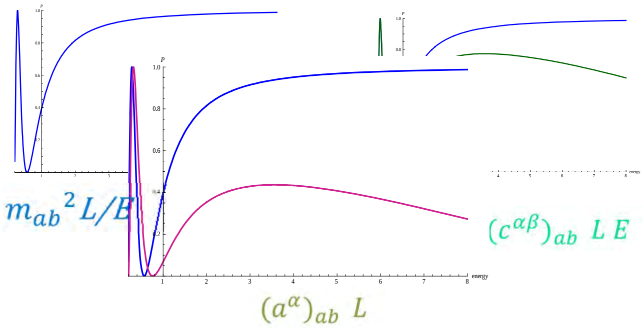
$$\sin^2(\Delta_{ab}\ L/2)$$

$m_{ab}{}^2 L/E$

$(a^\alpha)_{ab}\ L$

$(c^{\alpha\beta})_{ab}\ L\ E$

$$P(\nu_\mu \rightarrow \nu_\mu) = 1 - \sin^2(2\theta)\sin^2(1.27\Delta m^2 L/E)$$



Neutrinos,
In and Beyond the Standard Model:

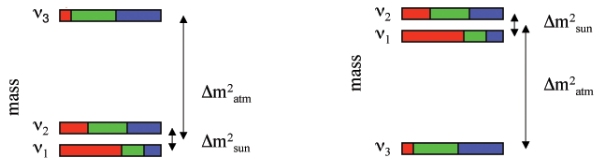
NEUTRINO MASS:

$\delta m^2_{atm} = 2.7^{+0.4}_{-0.3} \times 10^{-3} eV^2$ $\quad L/E = 500\ km/GeV$

$\delta m^2_{solar} = 8.0 \pm 0.4 \times 10^{-5} eV^2$ $\quad L/E = 15\ km/MeV$

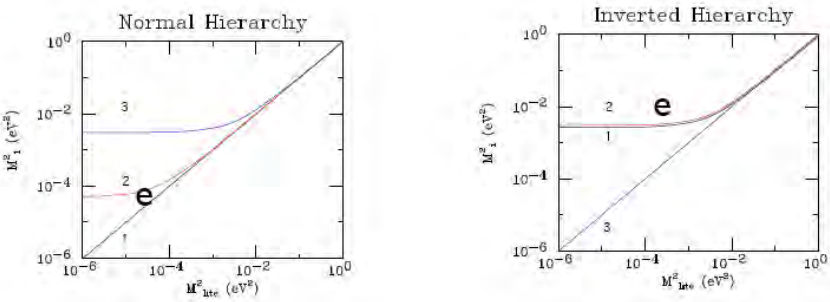
$$\Downarrow$$

$$m_\nu^{Heavy} > \sqrt{\delta m^2_{atm}} = 50 meV$$

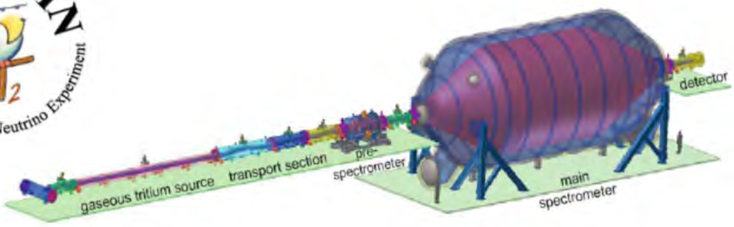


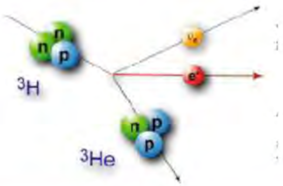
Normal mass hierarchy   Inverted mass hierarchy

Masses:



States 1 and 2 are ν_e rich.

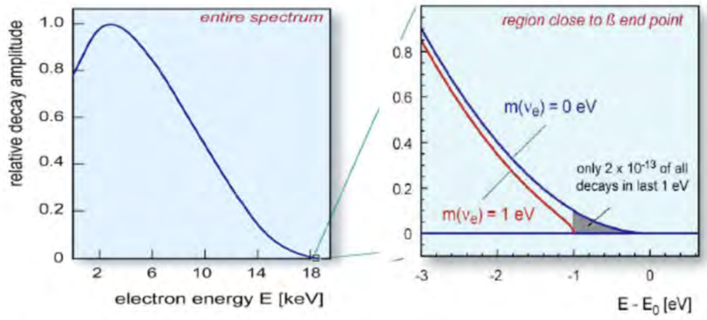




Requirements:
- Strong source
- Excellent energy resolution
- Small endpoint energy E_0
- Long term stability
- Low background rate

KATRIN Task:
Investigate Tritium endpoint with sub-eV precision

KATRIN Aim:
Improve m_ν sensitivity 10 x (2eV → 0.2eV)

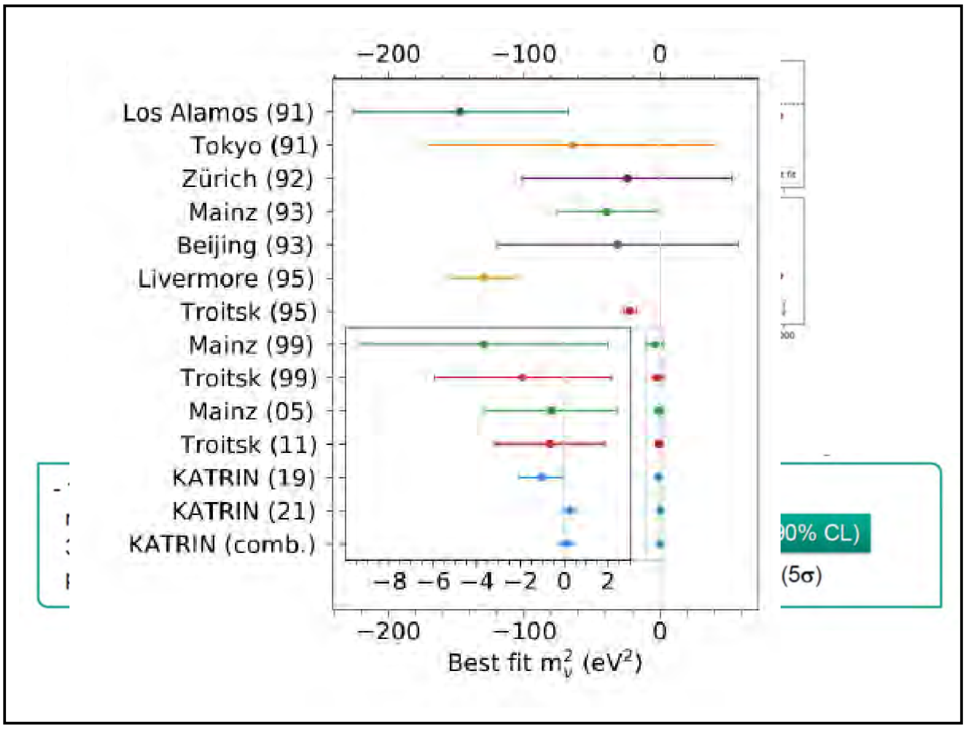


Decay Rate:

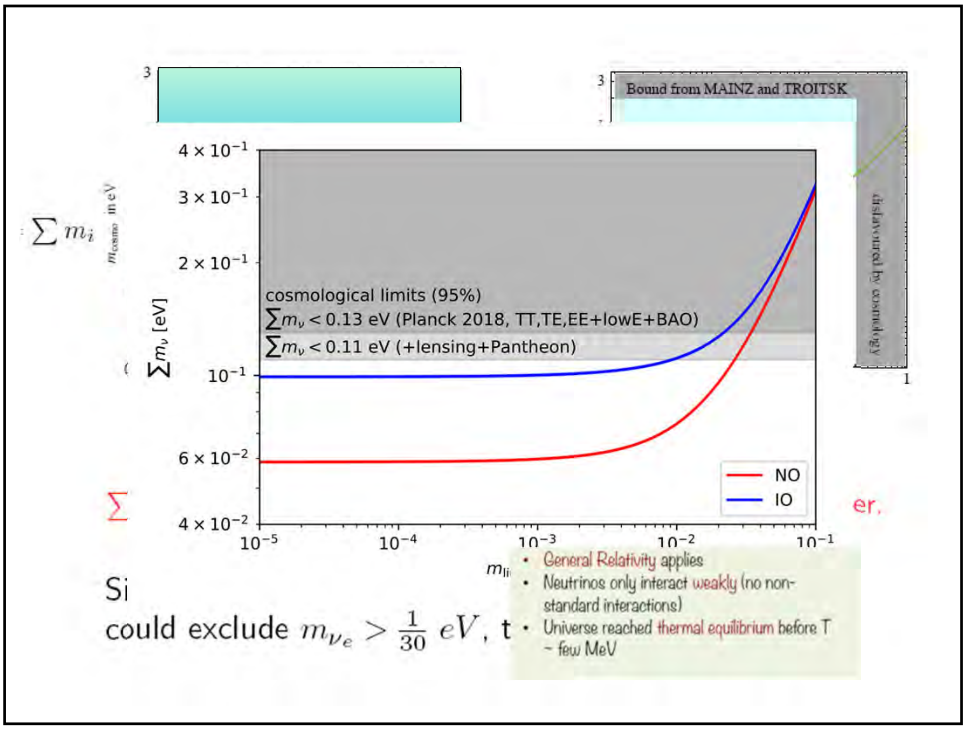
$$|\langle {}^3He + e^- + \bar{\nu}|T|{}^3H\rangle|^2 \sim pE(E_0 - E)\sum_k |U_{ek}|^2\sqrt{(E_0 - E)^2 - m_k^2}$$

if ν's quasi-degenerate: $m_1 \approx m_2 \approx m_3$

$$|\langle {}^3He + e^- + \bar{\nu}|T|{}^3H\rangle|^2 \sim pE(E_0 - E)\sqrt{(E_0 - E)^2 - m_\nu^2}$$



Los Alamos (91)
Tokyo (91)
Zürich (92)
Mainz (93)
Beijing (93)
Livermore (95)
Troitsk (95)
Mainz (99)
Troitsk (99)
Mainz (05)
Troitsk (11)
KATRIN (19)
KATRIN (21)
KATRIN (comb.)
−200 −100 0
−8 −6 −4 −2 0 2
Best fit m_ν^2 (eV^2)
0% CL)
(5σ)



Bound from MAINZ and TROITSK
disfavoured by cosmology
$\sum m_i$
cosmological limits (95%)
$\sum m_\nu < 0.13$ eV (Planck 2018, TT,TE,EE+lowE+BAO)
$\sum m_\nu < 0.11$ eV (+lensing+Pantheon)
$\sum m_\nu$ [eV]
4×10^{-1}
3×10^{-1}
2×10^{-1}
10^{-1}
6×10^{-2}
4×10^{-2}
10^{-5} 10^{-4} 10^{-3} 10^{-2} 10^{-1}
NO
IO
er.
Si
could exclude $m_{\nu_e} > \frac{1}{30}\ eV$, t
• General Relativity applies
• Neutrinos only interact weakly (no non-standard interactions)
• Universe reached thermal equilibrium before T ~ few MeV

CMB: neutrino mass

Spherical harmonics decomposition:

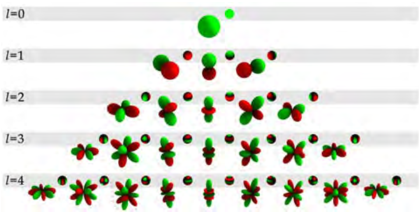
$$T(\hat{n}) = \sum_{\ell=0}^{\ell_{\max}} \sum_{m=-\ell}^{\ell} a_{\ell m} Y_{\ell m}(\hat{n})$$

With expansion coefficients:

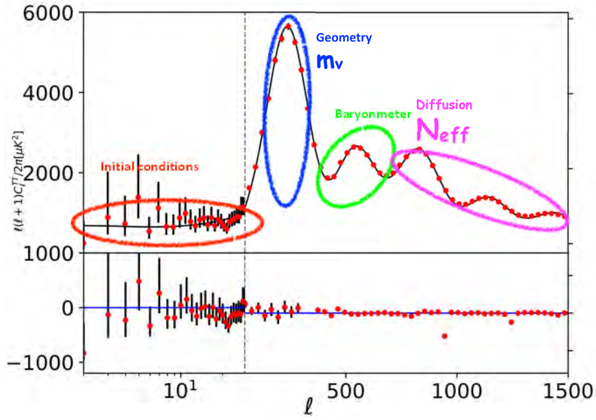
$$a_{\ell m} = \int_{4\pi} T(\hat{n}) Y^*_{\ell m}(\hat{n}) d\Omega$$

The angular power spectrum measures the amplitude of the expansion coefficients as a function of the wavelength:

$$C_\ell = \frac{1}{2\ell+1} \sum_{m=-\ell}^{\ell} |a_{\ell m}|^2$$



CMB: a lot to learn about….



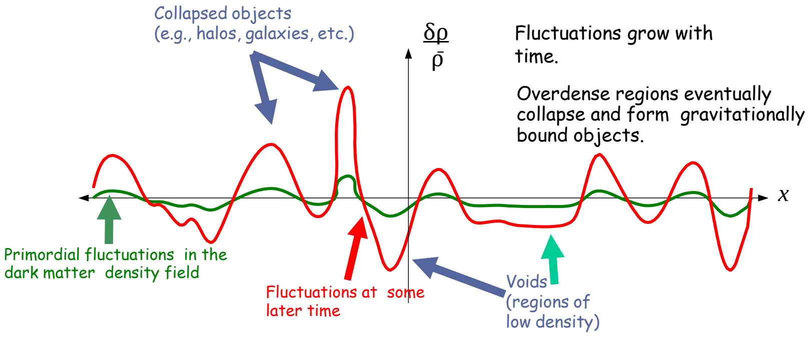
How structures form...

Photons freestream: Inhomogeneities turn into anisotropies

Initial fluctuations seeded by, e.g., inflation.

Fluctuations grow with time.

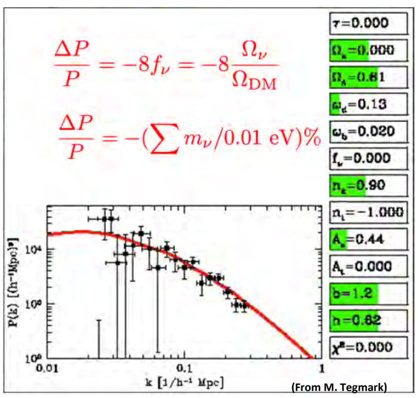
Overdense regions eventually collapse and form gravitationally bound objects.

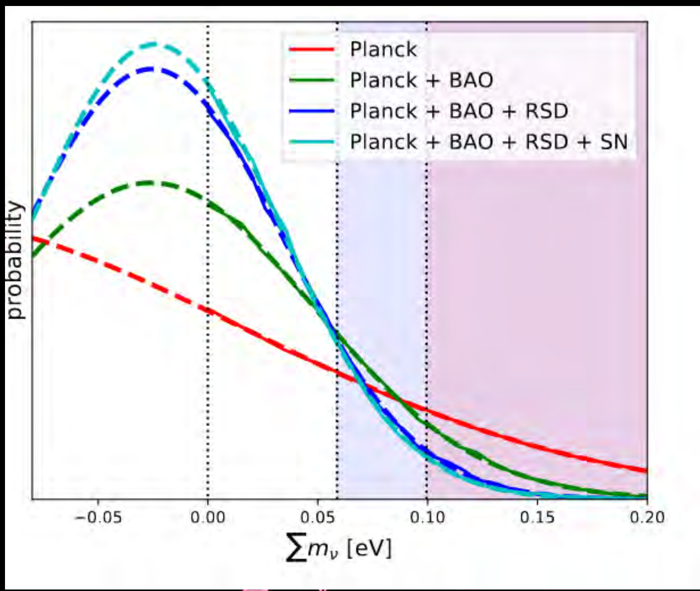


From Y. Wong

Large scale structure

Matter power spectrum suppression



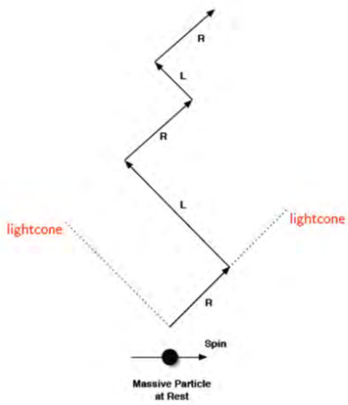


https://arxiv.org/2007.08991

+ more background probes

$\sum m_\nu <$ 0.043 eV at 95%CL https://arxiv.org/pdf/2405.03368

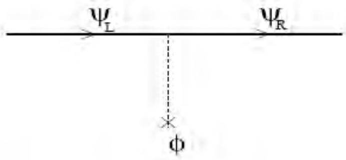
What is Fermion Mass ???



A mass can be thought of as a $L \leftrightarrow R$ transition:

$$m\,\overline{\psi_L}\psi_R + h.c.$$

In the SM fermion masses originate in the interaction with the Higgs field:



$$\lambda_f\,\overline{\psi_L}\,\Phi\,\psi_R + h.c. \rightarrow m_f = \lambda_f\,v$$

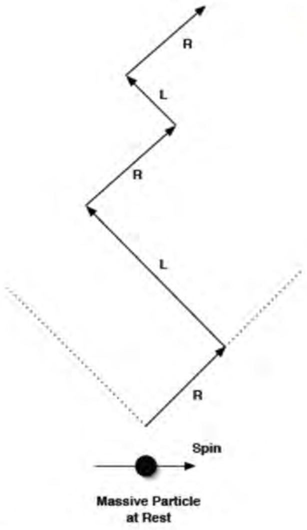
Fermion Masses:

	electron	positron	
Left Chiral	e_L	$\bar{e}_R$	SU(2)×U(1)
Right Chiral	e_R	$\bar{e}_L$	U(1)

CPT: $e_L \leftrightarrow \bar{e}_R$ and $e_R \leftrightarrow \bar{e}_L$

Mass couples L to R:

e_L to e_R AND also $\bar{e}_R$ to $\bar{e}_L$ Dirac Mass terms.



Mass couples L to R:

$$P^2 = M^2, \quad P \cdot S = 0 \quad and \quad S^2 = -1$$

$$u(P,S) = \frac{(1+\gamma_5)}{2} u(\frac{P+MS}{2}) + e^{i\phi} \frac{(1-\gamma_5)}{2} u(\frac{P-MS}{2})$$

right massless   left massless

A coupling of
e_L to $\bar{e}_R$ OR e_R to $\bar{e}_L$ would be (Majorana) mass term
but this violates conservation of electric charge!

Seesaw / Dirac Neutrinos / Light Sterile Neutrinos

	Nu	CPT:	Anti-Nu	
Left Chiral	ν_L	$\Leftrightarrow$	$\bar{\nu}_R$	
	$\Updownarrow$		$\Updownarrow$	Dirac Masses
Right Chiral	ν_R	$\Leftrightarrow$	$\bar{\nu}_L$	
		Majorana Masses		

Coupling of

- ν_L to ν_R AND $\bar{\nu}_R$ to $\bar{\nu}_L$ are the Dirac masses.
- ν_L to $\bar{\nu}_R$ forbidden by weak isospin.
- ν_R to $\bar{\nu}_L$ allowed and coefficient is unprotected. ($\rightarrow M$)

ν_L to $\bar{\nu}_R$ ν_L to ν_R

$$\begin{pmatrix} 0 & m_D \\ m_D & M \end{pmatrix}$$

$\bar{\nu}_R$ to $\bar{\nu}_L$ ν_R to $\bar{\nu}_L$

Two Majorana neutrinos
with masses m_D^2/M and M

Seesaw:
Yanagida, Gell-man-
Ramond-Slansky

- Coupling of ν_R to $\bar{\nu}_L$ allowed and coefficient is unprotected. ($\rightarrow M$)

Also applies to sterile neutrinos.

Light Sterile Neutrinos and/or Dirac Neutrinos Unexpected!!!

The consequences of this alternative are profound:

- Physics beyond the SM at a scale M!
- Majorana fermions carry no conserved charge: L is violated !

$$\nu_L \rightarrow e^{i\alpha}\nu_L$$

does not leave the Majorana mass term invariant.
$\rightarrow$ Most welcome for baryogenesis: a mechanism to understand the matter-antimatter asymmetry in the Universe emerges naturally
$\rightarrow$ Most welcome by string theory: it is difficult to get global $U(1)$ charges conserved

GB1

Leptogenesis

Baryon Asymmetry is created by a Lepton Asymmetry produced by the decays of super heavy Majorana Neutrinos.

$$\frac{\Gamma(N \to l^+ \phi^-) - \Gamma(N \to l^- \phi^+)}{\Gamma(N \to l^+ \phi^-) + \Gamma(N \to l^- \phi^+)}$$

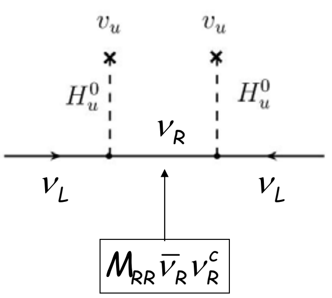
$\Gamma(N \to l^\pm \phi^\mp)$ depends on the Majorana Phases in the MNS mixing matrix.

$$B_{now} = \tfrac{1}{2}(B-L) + \tfrac{1}{2}(B+L) = \tfrac{1}{2}(B-L)_{ini} = -\tfrac{1}{2}L_{ini}$$

$\tfrac{1}{2}(B+L) \to 0$

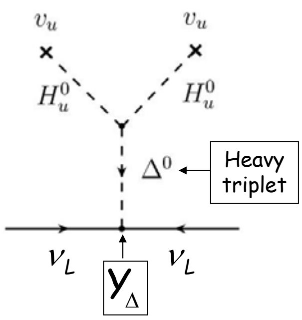
Types of see-saw mechanism

Type I see-saw mechanism



$$m_{LL}^{I} \approx -m_{LR} M_{RR}^{-1} m_{LR}^{T}$$

Type II see-saw mechanism



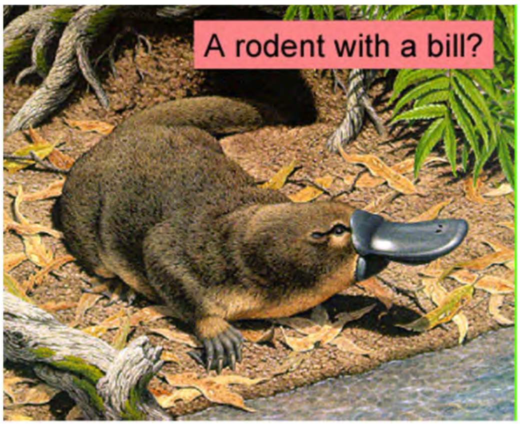
$$m_{LL}^{II}\, \bar{\nu}_L \nu_L^c \approx Y_\Delta \frac{v_u^2}{M_\Delta}$$

Diapositiva 41

GB1 Gabriela Barenboim; 29/04/2024

Naturalness may be over rated …

Does this look natural ??

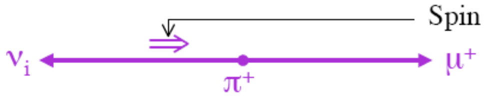


How Can We Demonstrate That $\overline{\nu}_i = \nu_i$?

We assume neutrino interactions are correctly described by the SM. Then the interactions conserve L ($\nu \rightarrow \ell^-$; $\overline{\nu} \rightarrow \ell^+$).

An Idea that Does Not Work
[and illustrates why most ideas do not work]

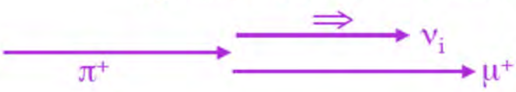
Produce a ν_i via—



Pion Rest Frame

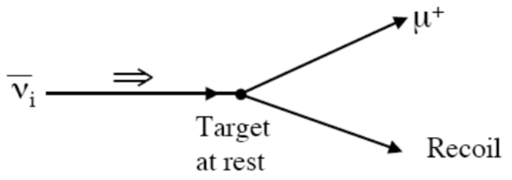
Give the neutrino a Boost:

$\beta_\pi(\text{Lab}) > \beta_\nu(\pi \text{ Rest Frame})$



Lab. Frame

The SM weak interaction causes—



$\nu_i = \overline{\nu}_i$ means that $\nu_i(h) = \overline{\nu}_i(h)$.

helicity

If $\nu_i \xrightarrow{\Rightarrow} \; = \; \overline{\nu}_i \xrightarrow{\Rightarrow}$,

our $\nu_i \xrightarrow{\Rightarrow}$ will make μ^+ too.

Minor Technical Difficulties

$$\beta_\pi(\text{Lab}) > \beta_\nu(\pi \text{ Rest Frame})$$

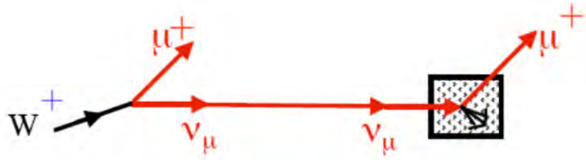
$$\Rightarrow \frac{E_\pi(\text{Lab})}{m_\pi} > \frac{E_\nu(\pi \text{ Rest Frame})}{m_{\nu_i}}$$

$$\Rightarrow E_\pi \text{ (Lab)} > 10^4 \text{ TeV} \quad \text{if } m_\nu \sim 1 \text{ eV}$$

Fraction of all π–decay that get helicity flipped

$$\approx \left(\frac{m_\nu}{E_\nu(\pi \text{ Rest Frame})} \right)^2 \sim 10^{-16} \quad \text{if } m_\nu \sim 1 \text{ eV}$$

For Majorana Neutrinos

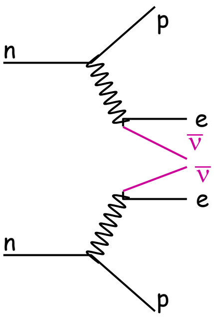


Not Observed

Allowed

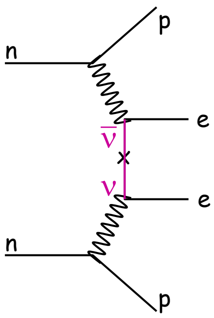
BUT Suppressed by $\frac{m_\nu^2}{E^2} \sim 10^{-20}$!!!

➢ How we can find out ?



SM double weak process

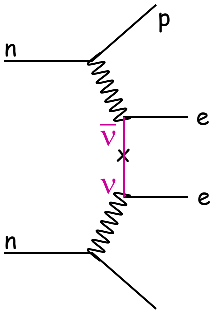
4 body decay: continuos spectrum for the e energy sum



Only allowed for Majorana ν

2 body decay: e energy sum is a delta

$\overline{\nu}_i$ is emitted (RH + O(m_i/E) LH)



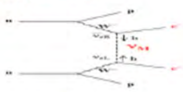
Amp[ν_i contribution] ~ m_i

$$\text{Amp}[0\nu\beta\beta] \propto \left| \sum m_i U_{ei}^2 \right|$$

effective mass

Neutrinoless double beta decay

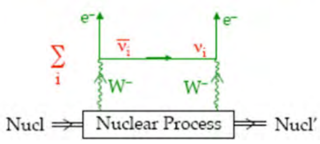
- Most sensitive (terrestrial) probe of the absolute neutrino mass
- Unique way of proving Majorana nature of ν
- If Majorana ν is the only mechanism, ===>



$$< m >_{\beta\beta} \equiv \left| \sum_{i=1}^{3} m_i U_{ei}^2 \right|$$

$$= \left| m_1 c_{12}^2 c_{13}^2 + m_2 s_{12}^2 c_{13}^2 e^{2i\beta} + m_3 s_{13}^2 e^{2i(\gamma-\delta)} \right|$$

$$T_{1/2}^{0\nu} = \frac{m_e}{G_{0\nu} M_{\rm nucl}^2 m_{ee}^2}$$



Best bounds from

^{136}Xe (KamLAND-ZEN):

$$T_{1/2}^{0\nu,\mathrm{Xe}} > 2.3 \times 10^{26} \text{ yr at 90\%CL}$$

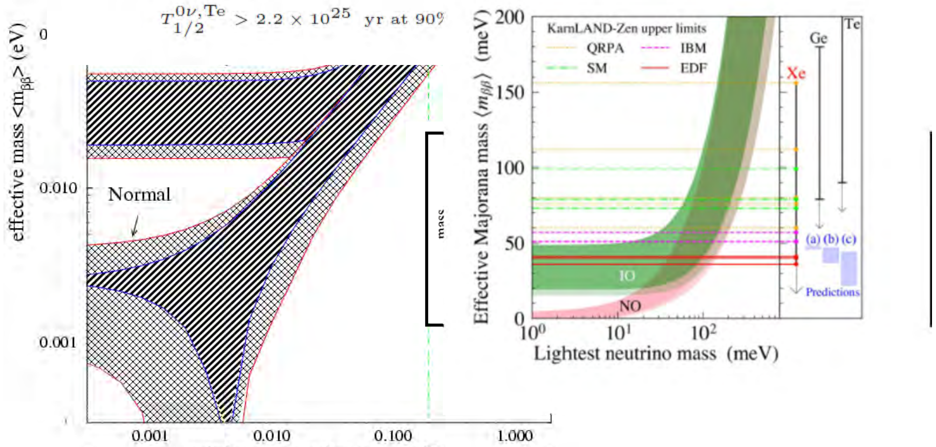
^{76}Ge (Gerda):

$$T_{1/2}^{0\nu,\mathrm{Ge}} > 1.8 \times 10^{26} \text{ yr at 90\%CL}$$

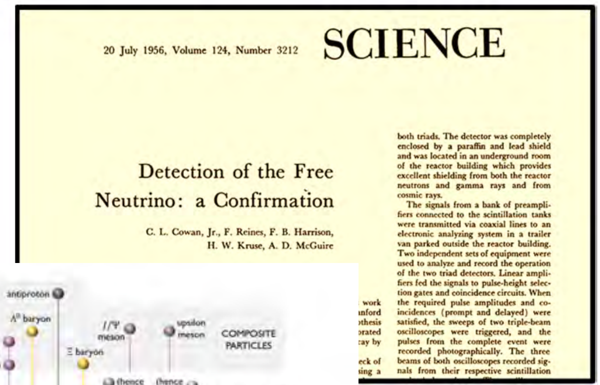
^{130}Te (Cuore):

$$T_{1/2}^{0\nu,\mathrm{Te}} > 2.2 \times 10^{25} \text{ yr at 90\%}$$

$$m_{\beta\beta} = \left| \sum m_i U_{ei}^2 \right|$$



NEUTRINOS AGE WELL



20 July 1956, Volume 124, Number 3212

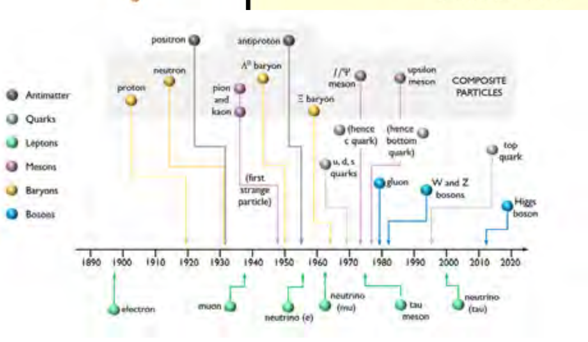
SCIENCE

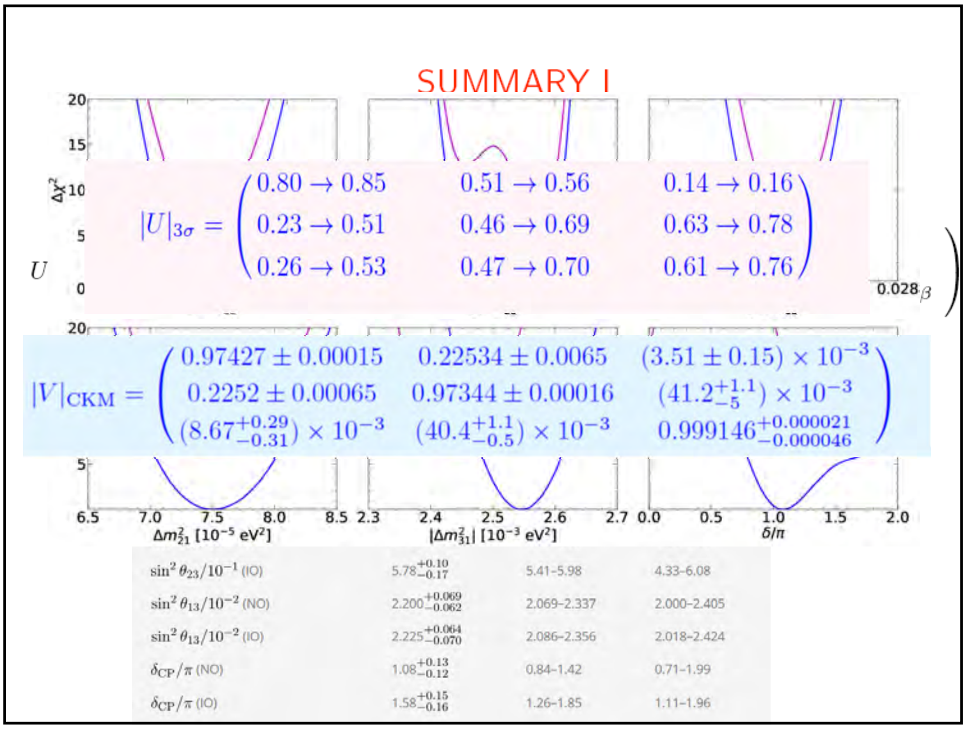
Detection of the Free Neutrino: a Confirmation

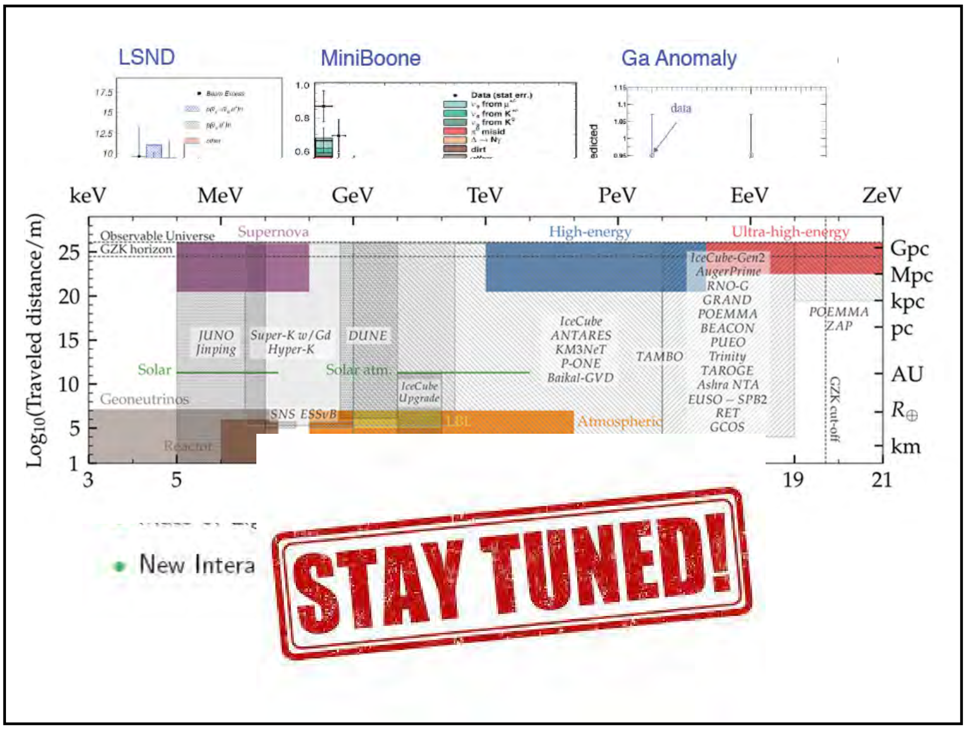
C. L. Cowan, Jr., F. Reines, F. B. Harrison, H. W. Kruse, A. D. McGuire

both triads. The detector was completely enclosed by a paraffin and lead shield and was located in an underground room of the reactor building which provides excellent shielding from both the reactor neutrons and gamma rays and from cosmic rays.

The signals from a bank of preamplifiers connected to the scintillation tanks were transmitted via coaxial lines to an electronic analyzing system in a trailer van parked outside the reactor building. Two independent sets of equipment were used to analyze and record the operation of the two triad detectors. Linear amplifiers fed the signals to pulse-height selection gates and coincidence circuits. When the required pulse amplitudes and coincidences (prompt and delayed) were satisfied, the sweeps of two triple-beam oscilloscopes were triggered, and the pulses from the complete event were recorded photographically. The three beams of both oscilloscopes recorded signals from their respective scintillation




SUMMARY I
$|U|_{3\sigma} = \begin{pmatrix} 0.80 \to 0.85 & 0.51 \to 0.56 & 0.14 \to 0.16 \\ 0.23 \to 0.51 & 0.46 \to 0.69 & 0.63 \to 0.78 \\ 0.26 \to 0.53 & 0.47 \to 0.70 & 0.61 \to 0.76 \end{pmatrix}$
$|V|_{CKM} = \begin{pmatrix} 0.97427 \pm 0.00015 & 0.22534 \pm 0.0065 & (3.51 \pm 0.15) \times 10^{-3} \\ 0.2252 \pm 0.00065 & 0.97344 \pm 0.00016 & (41.2^{+1.1}_{-5}) \times 10^{-3} \\ (8.67^{+0.29}_{-0.31}) \times 10^{-3} & (40.4^{+1.1}_{-0.5}) \times 10^{-3} & 0.999146^{+0.000021}_{-0.000046} \end{pmatrix}$
Δm^2_{21} [10^{-5} eV^2]
$|\Delta m^2_{31}|$ [10^{-3} eV^2]
δ/π
$\Delta\chi^2$
$\sin^2\theta_{23}/10^{-1}$ (IO) $5.78^{+0.10}_{-0.17}$ 5.41–5.98 4.33–6.08
$\sin^2\theta_{13}/10^{-2}$ (NO) $2.200^{+0.069}_{-0.062}$ 2.069–2.337 2.000–2.405
$\sin^2\theta_{13}/10^{-2}$ (IO) $2.225^{+0.064}_{-0.070}$ 2.086–2.356 2.018–2.424
δ_{CP}/π (NO) $1.08^{+0.13}_{-0.12}$ 0.84–1.42 0.71–1.99
δ_{CP}/π (IO) $1.58^{+0.15}_{-0.16}$ 1.26–1.85 1.11–1.96


LSND
MiniBoone
Ga Anomaly
keV MeV GeV TeV PeV EeV ZeV
Log10(Traveled distance/m)
Observable Universe
GZK horizon
Supernova
High-energy
Ultra-high-energy
Gpc Mpc kpc pc AU $R_\oplus$ km
JUNO Jinping
Super-K w/Gd Hyper-K
DUNE
IceCube ANTARES KM3NeT P-ONE Baikal-GVD
TAMBO
IceCube-Gen2 AugerPrime RNO-G GRAND POEMMA BEACON PUEO Trinity TAROGE Ashra NTA EUSO – SPB2 RET GCOS
POEMMA ZAP
Solar
Solar atm.
IceCube Upgrade
Geoneutrinos
SNS ESSνB
LBL
Atmospheric
Reactor
GZK cut-off
New Intera
STAY TUNED!